

Introduction to Particle Physics

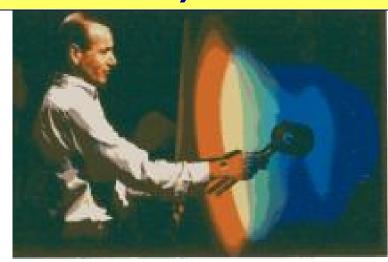
(for non physics students)

4. UNIFIED UNIVERSE

(no strings attached)

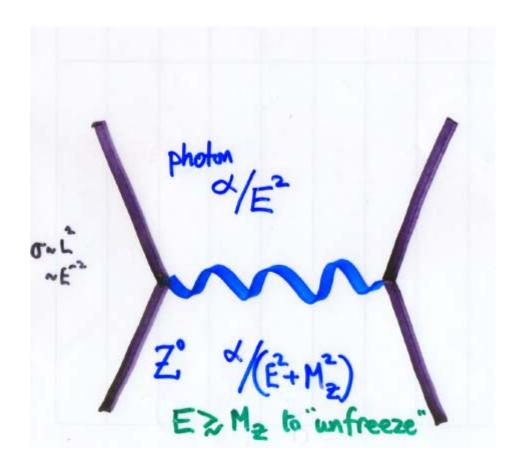


PROFESSOR FRANK CLOSE
EXETER COLLEGE
UNIVERSITY OF OXFORD



FORCES Summary (remember that waves -> particles)

NAME	action	CARRIER
Gravity	keeps us on ground	graviton?
Electromag	netic electrons in adoms Solids Stops us falling to centre of Earth	byoton (8)
Weak	β-ractioactivity P→He in Sun	W+ W- Z0
Strong	quarks glued inside p.n P.n in nuclei	gluons (g) 8 different
Only the	weak force carriers	have MASSes
My ~	BO GeV/c²	
	91 GeV/c2	



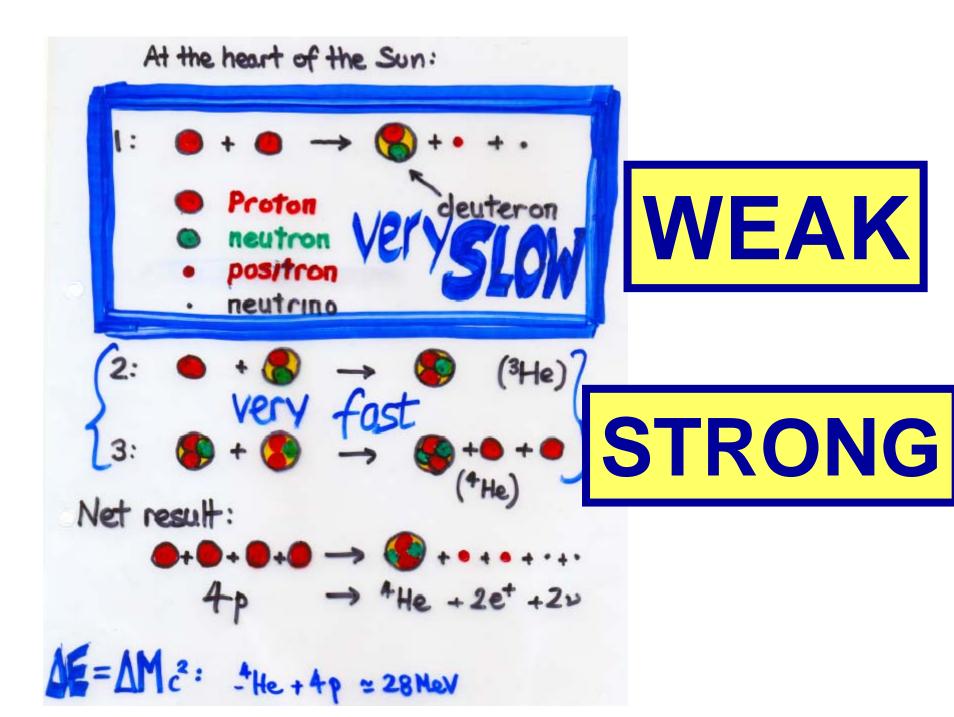
Feynman rules:

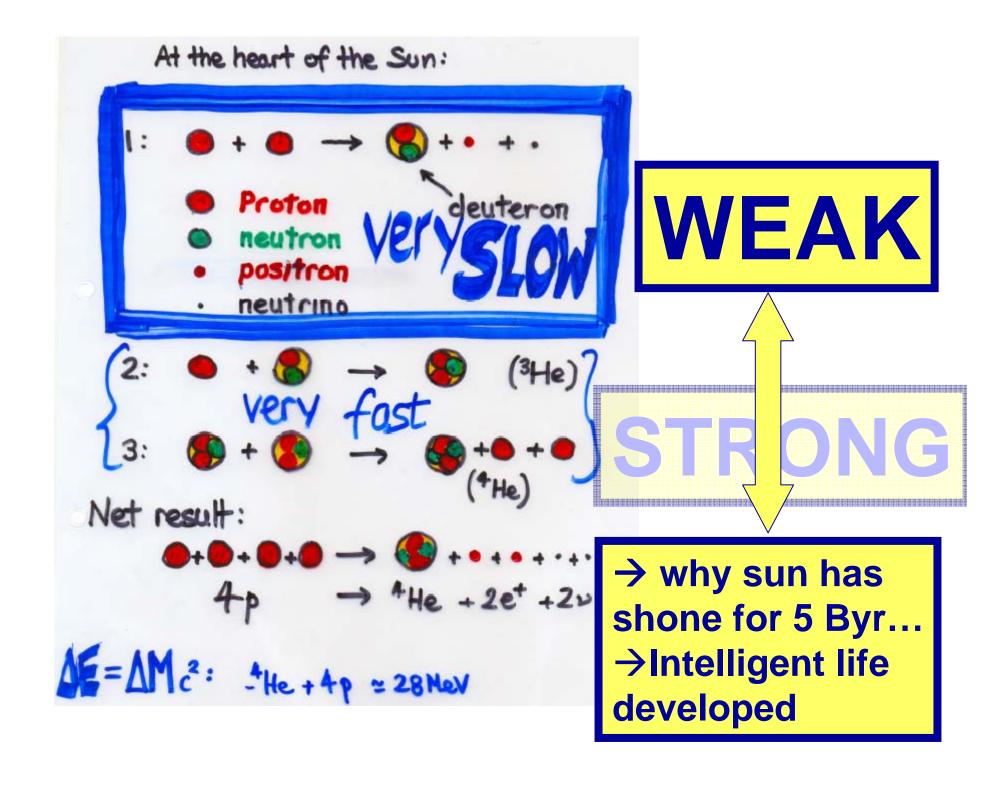
If energy E flows through the transmitted "virtual" particle (photon; Z) it costs 1/(E^2+M^2)



If E >> M the cost is 1/E^2....like the case of the photon

Only appears weak at low energy. Unified at high energy



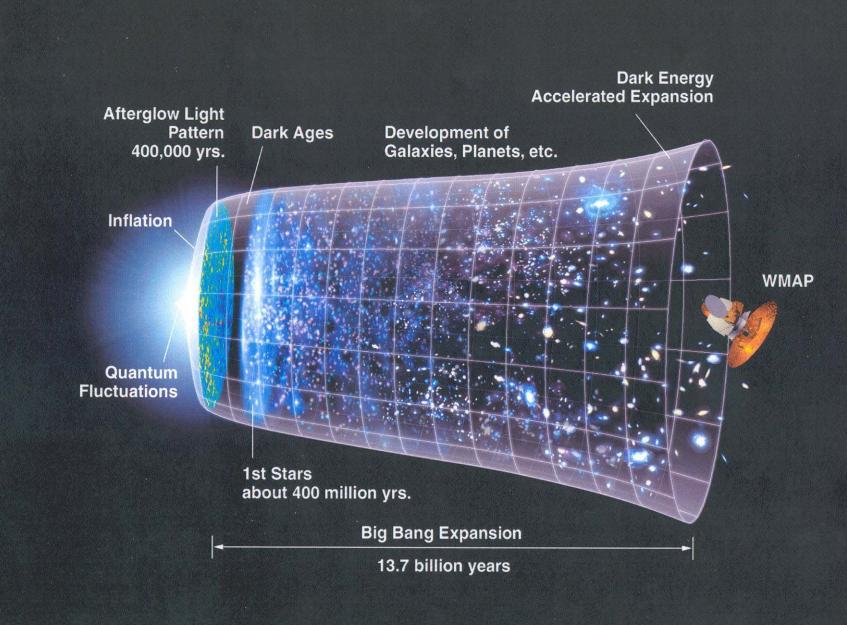


The weak force is feeble in the Sun ...

..because 10,000,000K ~ 1 keV << 80 GeV

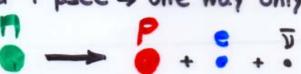
...this is why the sun has stayed active long enough for us to have evolved and be having this conversation.

- →We exist because m(W) is not zero
- → Mass matters





Temperature (energy) drops >>
After 1 usec >> one way only:



But at the same time:



then like processes 2 and 3 in the Sun until all the neutrons have gone

* MAKING

or

particles so far apart in the expanding universe that they no longer interact

T = 1 usec after BIG BANG

("the first fossils in the Universe)

move at high speed and if they have mass they start clustering together →contribute to formation of galaxies

Billion > per atom

if m(v) > m(proton)/109 = lev

they will dominate mass
density of the Universe

of future of universe and its formation

Universe expands - and cools expansion rate Rate depends on pressure

Rate depends on pressure which depends on the temperature in the gas and the number of neutrinos inside the gas volume (density) and this # depends on number of neutrino species

T=3 minutes after BIG BANG

75% protons
24%. Helium Nuclei
+ small amount of deuterons
+ free electrons.

Helium abundance + throngs of other

depends on expansion rate of the Universe which depends on number of neutrino species

Deuterium abundance depends on density of "ordinary matter" in the Universe.

T=3 minutes after BIG BANG

75% protons
24%. Helium Nuclei
+ small amount of deuterons
+ free electrons.

Helium abundance + trocas of other

depends on expansion rate of the Universe which depends on number of neutrino species

Deuterium abundance depends on density of "ordinary matter" in the Universe.

T=3 minutes after BIG BANG

75% protons
24% Helium Nuclei
+ small amount of deuterons
+ free electrons.

Action abundance*: House of other depends on expenses and the seconds

depends on expansion rate of the Universe which depends on number of neutrino species

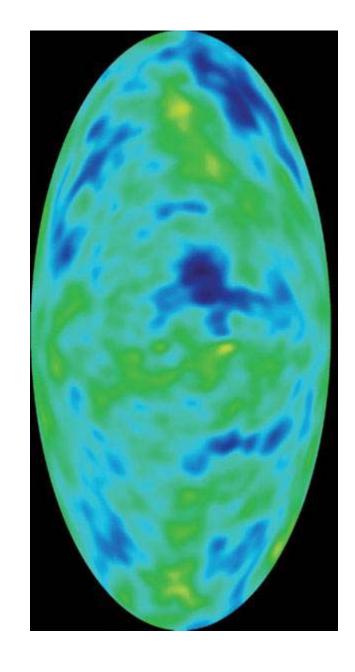
Deuterium abundance

depends on density of "ardinary matter" in the Universe.

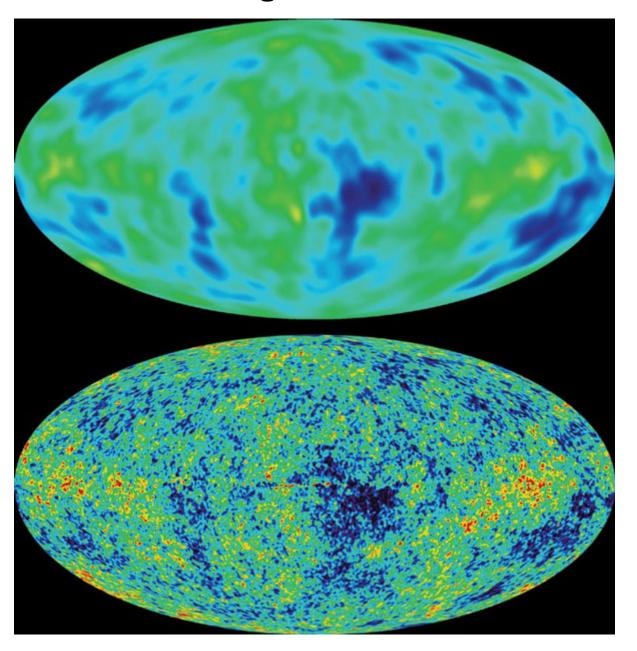
density of ordinary matter 2< total in universe

part of DARK MATTER puzzle

Time Passes. Temp drops 300,000 years later Exlock electrons combine with nuclei and make neutral adoms electromagnetic radiation was set free Universe becomes transparent 10 years later Emag 2 stretched: Microniane Rand. Black body background 3K (small fluctuations in Manuare rad = hints of proto structures, galaxier in early universe)



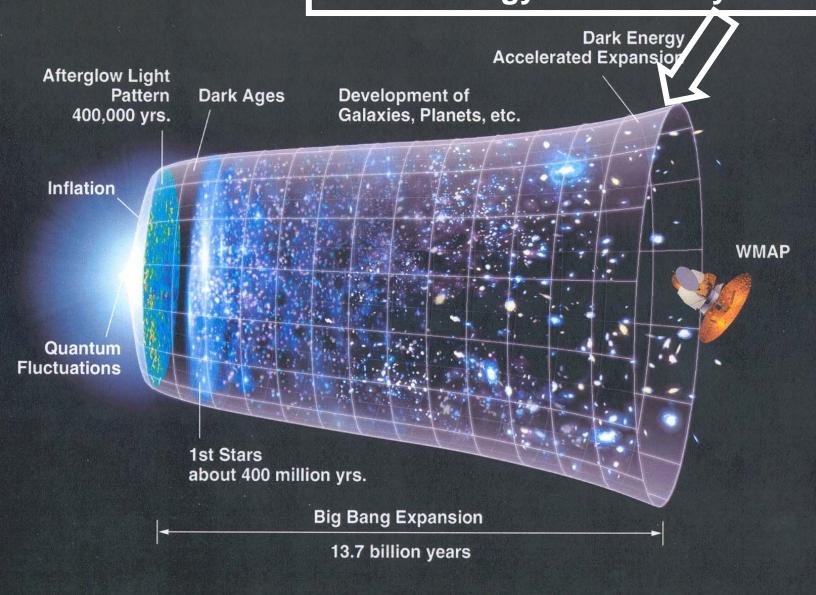
3K microwave bgnd now seen to have structure



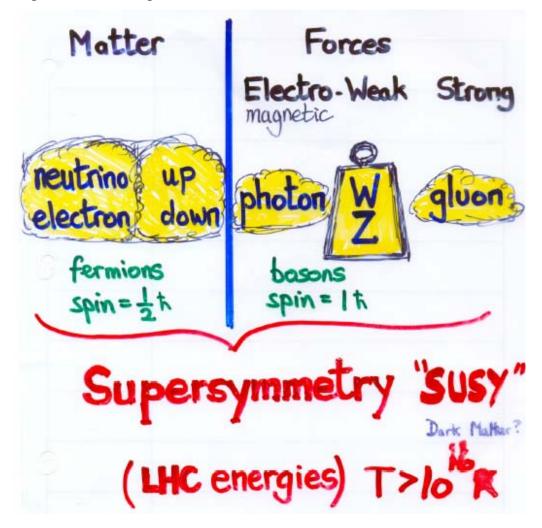
COBE 2000

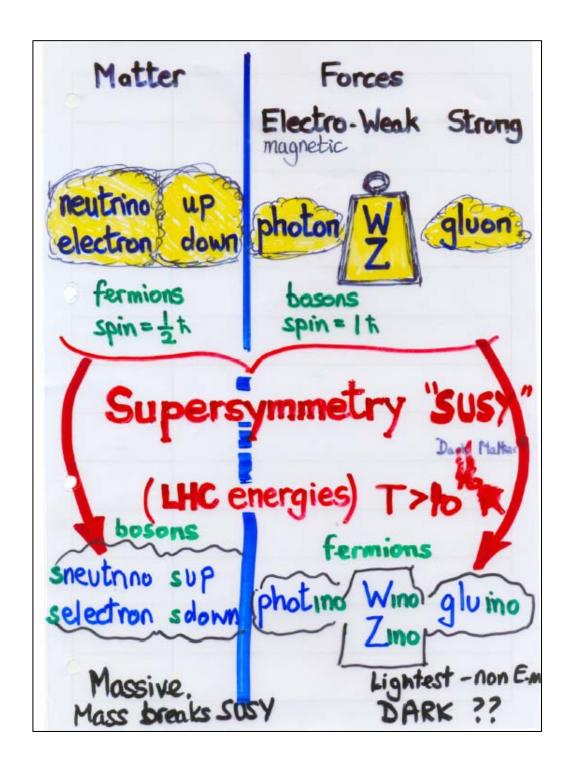
WMAP 2003

5byr ago accelerated expansion = Dark Energy. What? Why?.....



One further symmetry??





Whole new families to be found

Particle Physics @ CERN. Standard Model of PHes + Forces. · Quarks + Leptons. Spin 1/2 fermions · & W = gluons Spin 1 gauge bosons · Higgs Spin O boson Origins of matter. Structures + pallerns at E& ITEV Symmetry revealed at E > O(TeV) Forces (and particles) unified - SUSY. Some current big pazzles. . Dark Malter, Solar v, Massive v? (all the same?) . Why 3 generations What is difference between Mand M? I same? The Fifth Dimension

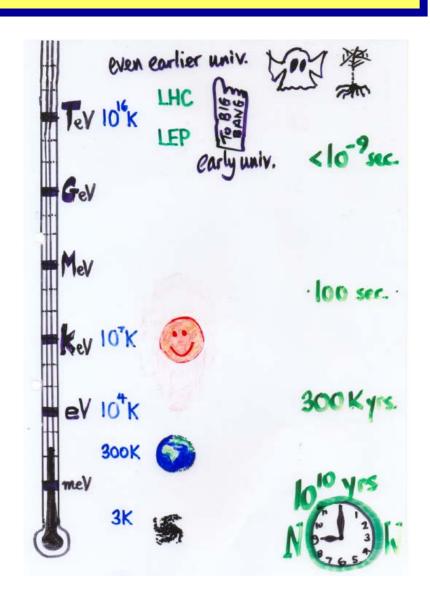
Finale: A glimpse of the future

recall from lecture 1.....

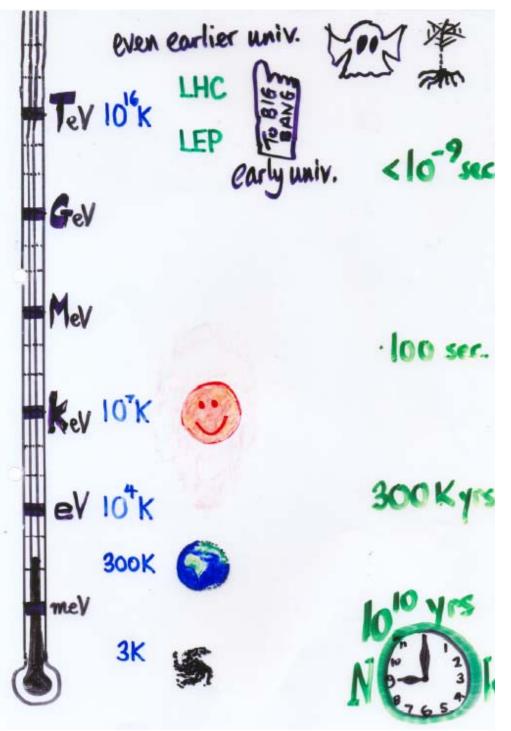
The Universe

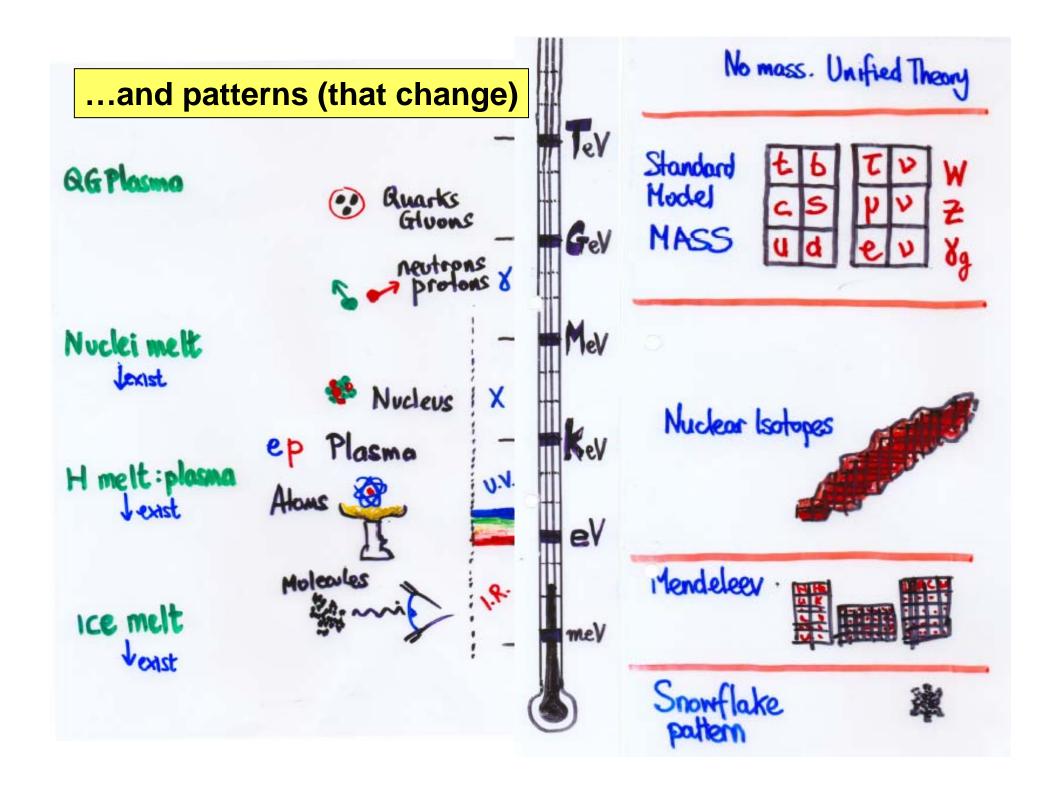
in

Temperature
Energy and
Time

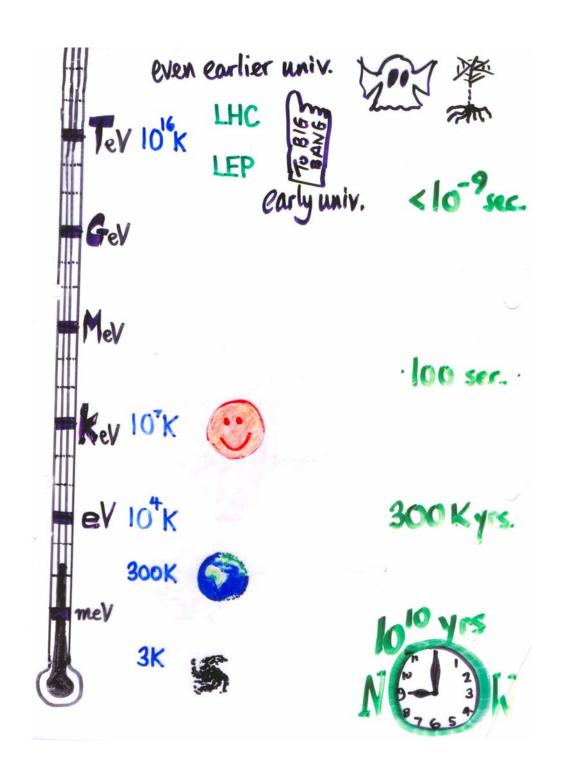


...and the nature of matter **QG Plasma** Auarks Gluons Nuclei melt Jexist Nucleus ** ep Plasma H melt:plasma V.V. Molecules ice melt Venst





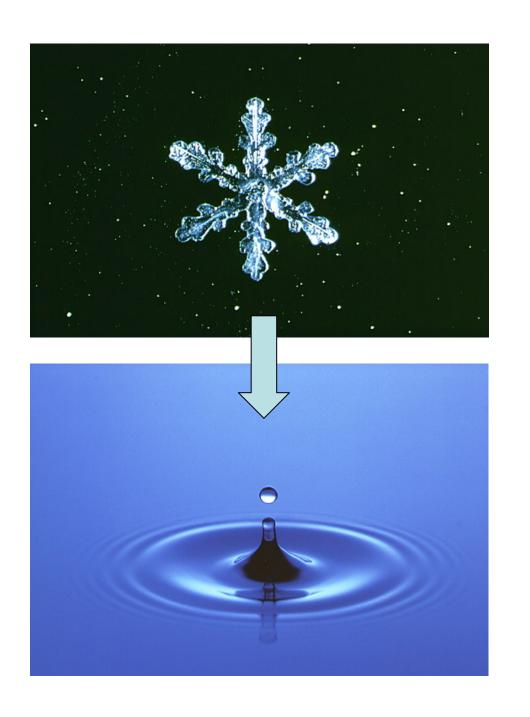
Temperature and symmetry in the universe



The Idea

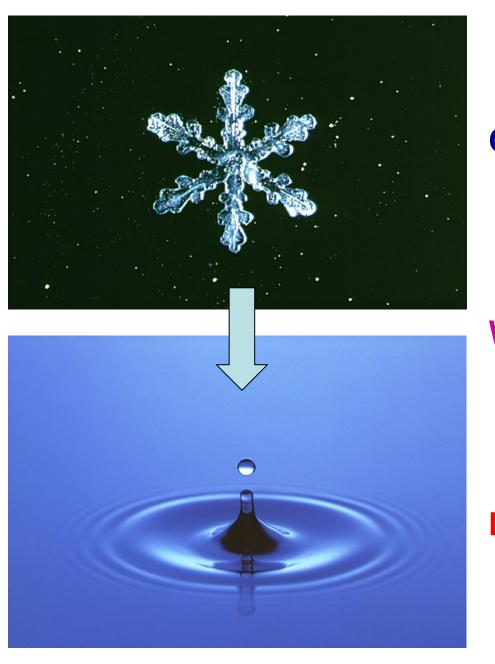


(I will tell you when to be cautious about inhaling)

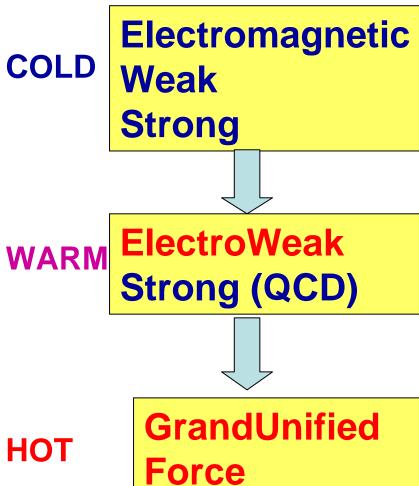


patterns and structures when cold (low energy)

Symmetry when warm (high energy)



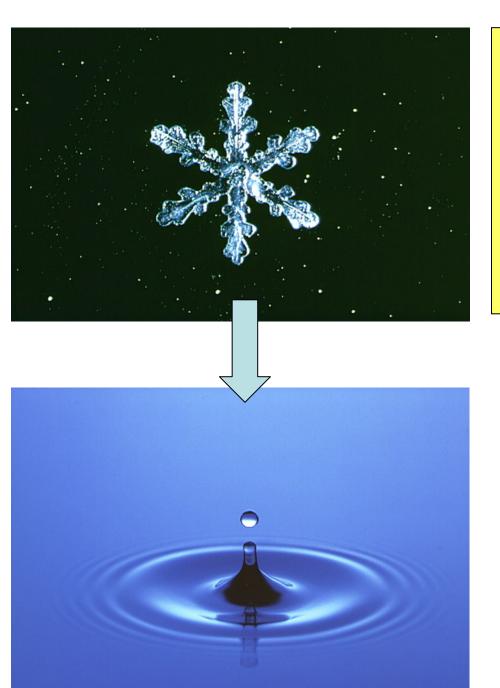
FORCES 1955-2005





Standard Model of Quarks Leptons and forces

- = pattern based on mass
- "cold" ="low" energy
- = below 1 TeV



Standard Model of Quarks Leptons and forces

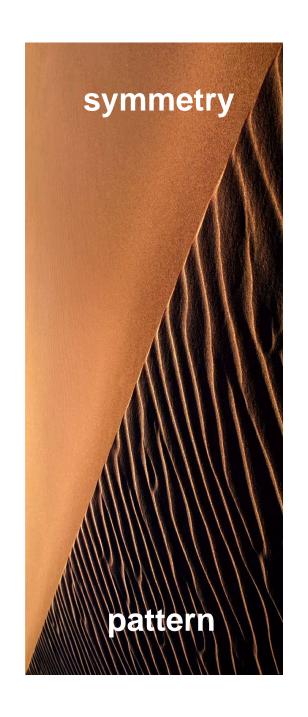
= pattern based on mass

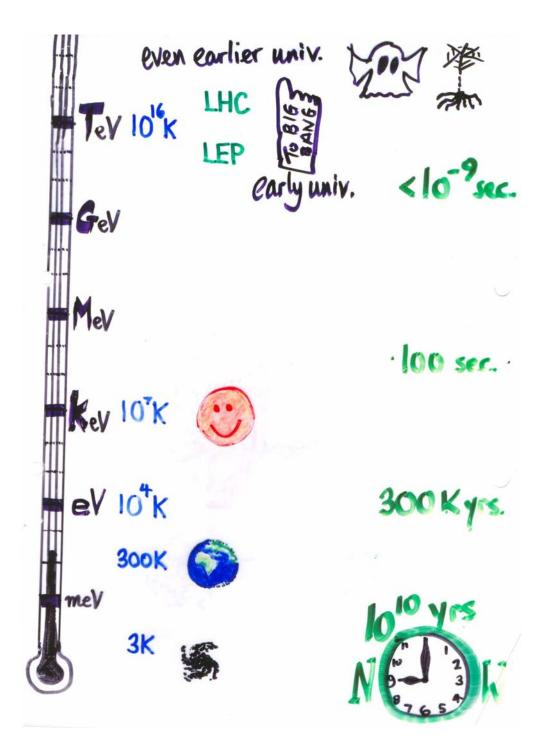
"cold" ="low" energy

= below 1 TeV

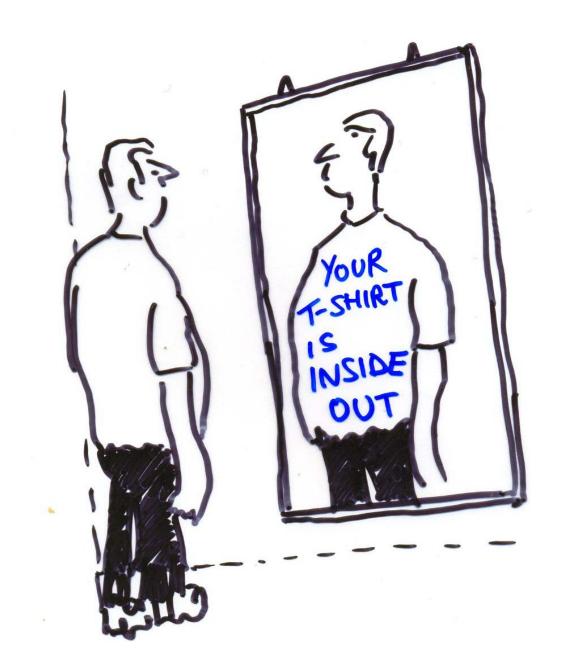
superSymmetry
when "warm"
(= high energy > 1TeV)

Higgs Boson Supersymmetry Nature of Reality

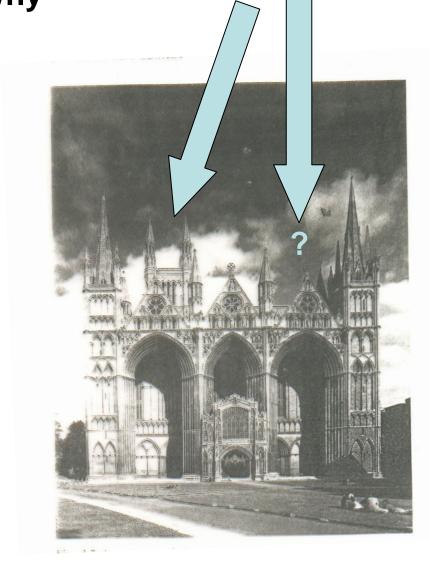


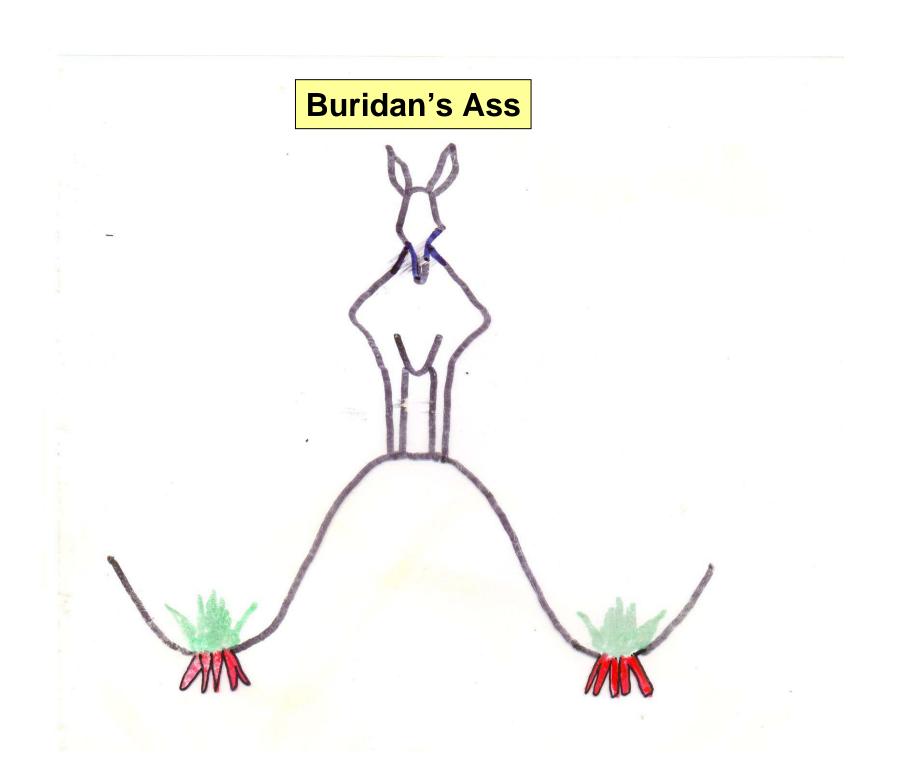


5. symmetries can disappear or change

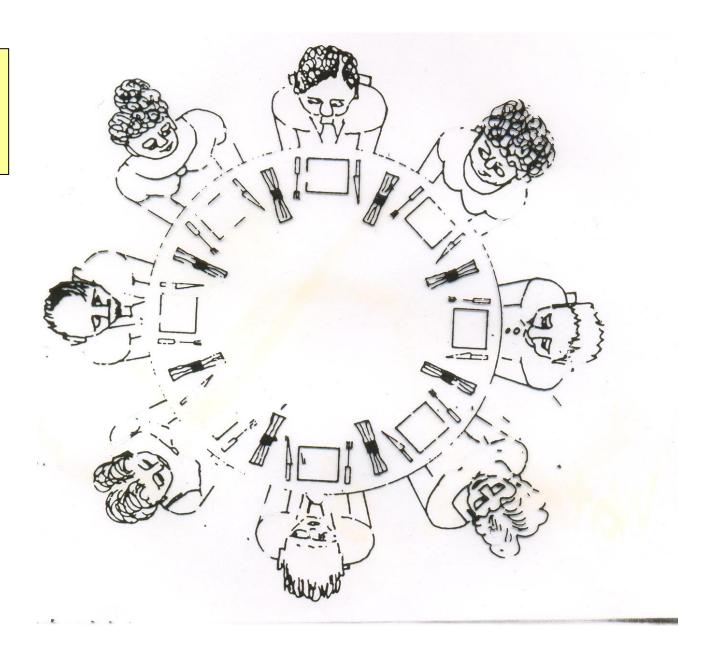


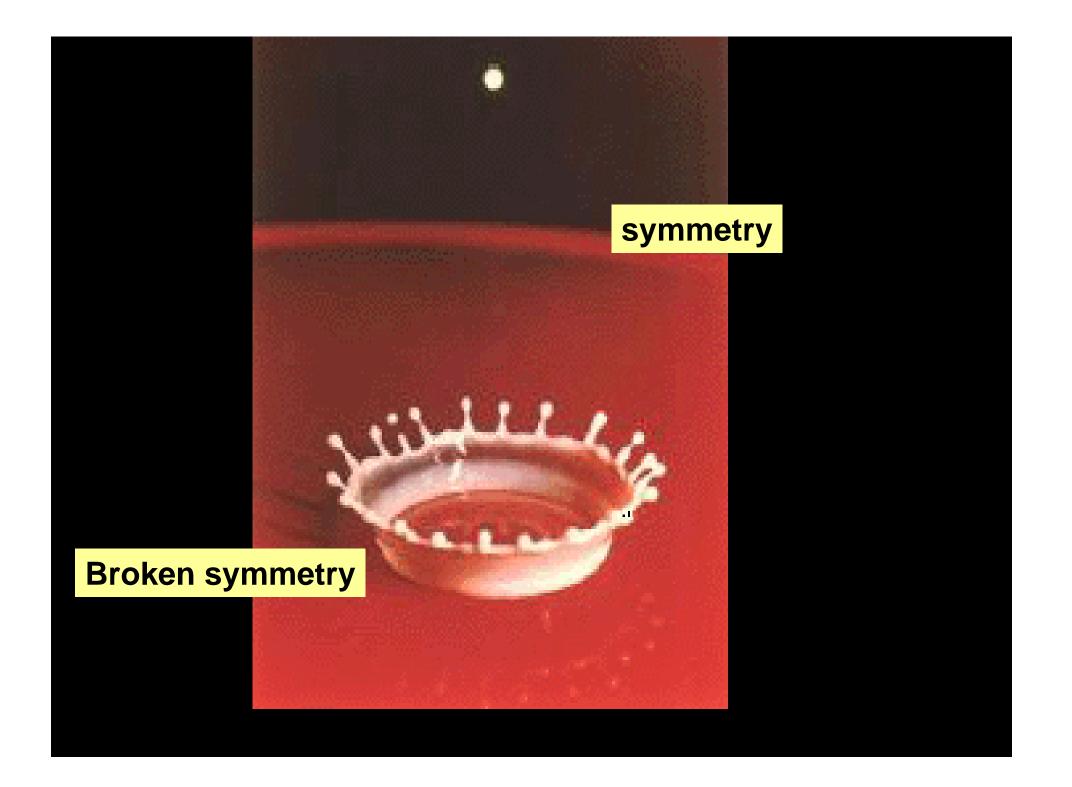
We like symmetry and when its absent we want to know why

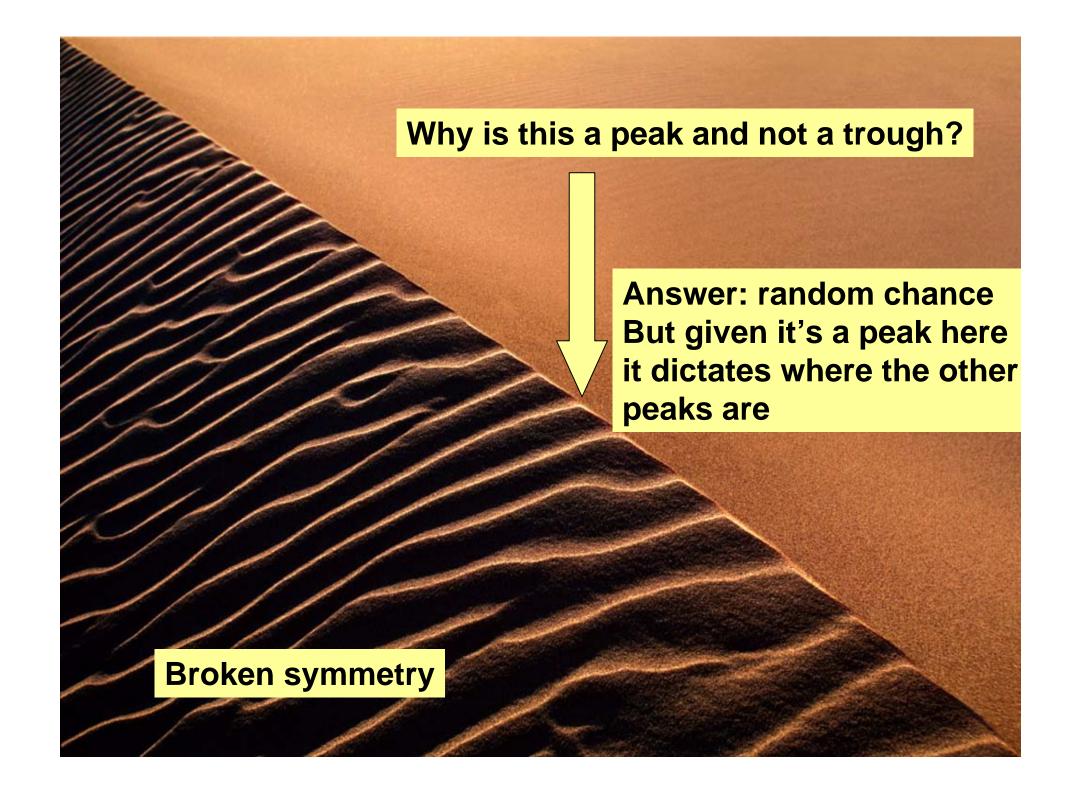




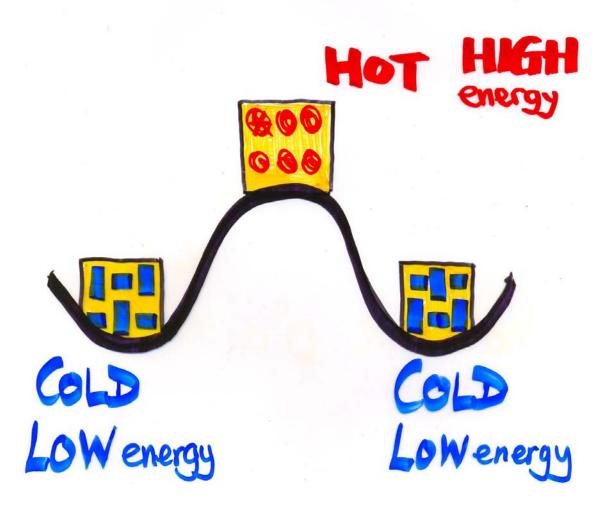
The problem of the symmetric dinner party

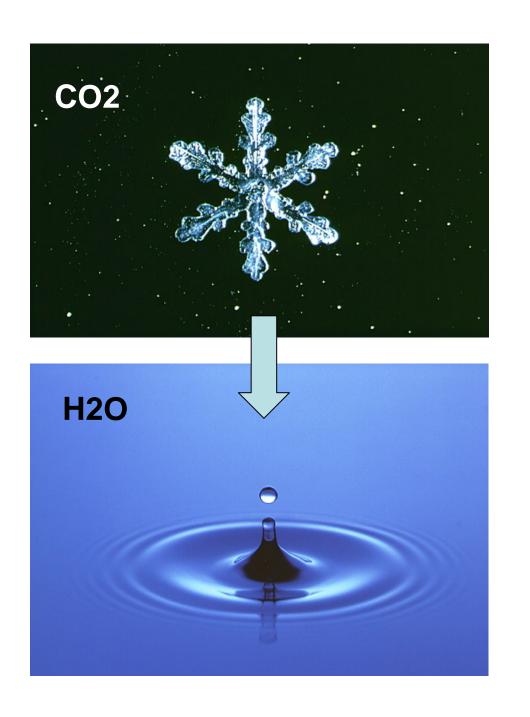






Magnets





patterns
and structures
when cold
(low energy)

Symmetry when warm (high energy)

As the universe cooled after the hot big bang.....
We think that an elegant symmetry......

... "froze" into structures And patterns

Such as Atoms Mendeleev's periodic table, And particles Quarks, forces and the Standard Model

... which is a pattern based on MASS

2008: heat up to energies above 1000 GeV = "1 TeV" and discover the origin of MASS (= Higgs?)

